



Forecasts on the contamination induced by unresolved point sources in primordial non-Gaussianity beyond *Planck*

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ABSTRACT

In this paper we present forecasts of the contamination on different shapes of the primordial non-Gaussianity f_{nl} parameter – detectable on future cosmic microwave background (CMB) high-resolution anisotropy maps – produced by unresolved extragalactic point sources at frequencies of cosmological interest (45–375 GHz). We consider two scenarios: an ideal (noiseless) mission and a possible future space-borne satellite, with instrumental characteristics similar to the ones proposed for the *Cosmic Origins Explorer (CORE)*. The local, equilateral, orthogonal and flat shapes are considered in both temperature (intensity) and polarized emission data. The angular power spectrum and bispectrum of extragalactic point sources are estimated by state-of-the-art models of source number counts. The impact of all the most relevant (far-infrared and radio selected) source populations on these shapes at *CORE* frequencies is studied. The results of this analysis show that unresolved extragalactic point sources should not induce a very relevant non-Gaussian signal in the frequency range 100–200 GHz, thus not preventing a correct estimate of the CMB primordial f_{nl} parameter. Polarization information allows one to significantly reduce the error-bars in the f_{nl} parameter and the bias induced by unresolved sources and, hence, to widen the range of frequencies for f_{nl} studies. On the contrary, at $\nu < 100$ GHz or $\nu > 225$ GHz, important non-Gaussian deviations in CMB anisotropy maps are expected due to unresolved extragalactic sources.

Key words: methods: data analysis – cosmic background radiation – radio continuum: galaxies – submillimetre: galaxies.

1 INTRODUCTION

The high precision achieved during the last years to detect possible weak non-Gaussian signals with primordial origin in the most precise cosmic microwave background (CMB) anisotropy data maps [such as *Wilkinson Microwave Anisotropy Probe (WMAP)*¹ data, the forthcoming *Planck*² data and future missions] stresses the importance of the characterization of all the possible non-primordial contaminants. The standard models of inflation (Starobinskiĭ 1979;

Guth 1981; Albrecht & Steinhardt 1982; Linde 1982, 1983; Mukhanov, Feldman & Brandenberger 1992) predict that the primordial fluctuations, linearly imprinted in the CMB anisotropies, are nearly compatible with a Gaussian distribution. Any departure from Gaussianity with a primordial origin would challenge our understanding of the physics of the early Universe. From a practical point of view, the first approach to detect non-Gaussian deviations is the study of the third-order moments (the so-called bispectrum) of the anisotropy map. For most of the cases with a physical interest, this bispectrum is usually characterized by a single amplitude, the non-linear coupling parameter f_{nl} (see e.g. Verde et al. 2000; Komatsu & Spergel 2001; Babich, Creminelli & Zaldarriaga 2004; Bartolo et al. 2004; Babich 2005). For recent reviews of the status of the primordial non-Gaussianity characterization see e.g. Bartolo,

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¹ <http://map.gsfc.nasa.gov/>

² <http://www.esa.int/planck>

Matarrese & Riotto (2010a), Yadav & Wandelt (2010), Liguori et al. (2010) and Martínez-González & Planck Collaboration (2012).

The best estimates of the f_{nl} parameter can currently be obtained using the publicly available *WMAP* data. Three different shapes with relevant interest for the physics of inflation have been constrained using these data: the local, equilateral and orthogonal f_{nl} shapes. Information about the inflationary scenarios that can produce these and other non-Gaussian shapes can be found for example in the reviews by Bartolo et al. (2004) and Yadav & Wandelt (2010). The instrumental noise contamination and the beam size of the *WMAP* 61 and 94 GHz frequency channels allow a maximum resolution of $\ell_{\text{max}} \sim 1000$. The best estimates of these shapes provided by the *WMAP* collaboration (Komatsu et al. 2011) are $\hat{f}_{\text{nl}}^{\text{loc}} = 32 \pm 21$, $\hat{f}_{\text{nl}}^{\text{eq}} = 26 \pm 140$ and $\hat{f}_{\text{nl}}^{\text{ort}} = -202 \pm 104$ (68 per cent CL) for the local, equilateral and the orthogonal shapes, respectively (see e.g. Curto et al. 2011; Curto, Martínez-González & Barreiro 2012, for independent analyses with similar results).

Unresolved point sources introduce non-Gaussian fluctuations in CMB anisotropies at a detectable level. In the past, their ‘contamination’ to the CMB angular bispectrum has been studied in detail by many authors (e.g. Refregier, Spergel & Herbig 2000; Komatsu & Spergel 2001; Argüeso, González-Nuevo & Toffolatti 2003; González-Nuevo, Toffolatti & Argüeso 2005; Argüeso et al. 2006; Babich & Pierpaoli 2008; Serra & Cooray 2008). As an example, the contamination of the f_{nl} parameter due to point sources in *WMAP* data has been studied by Komatsu et al. (2009) considering the simple constant-flux model to simulate the radio sources. They found that the bias introduced in the local and the equilateral shapes is $\Delta f_{\text{nl}}^{(\text{loc})} \sim 2$, $\Delta f_{\text{nl}}^{(\text{eq})} \sim 22$. Taking into account the source number counts model by de Zotti et al. (2005), Curto et al. (2011) estimated that the bias in f_{nl} is $\Delta f_{\text{nl}}^{(\text{loc})} \sim 2.5$, $\Delta f_{\text{nl}}^{(\text{eq})} \sim 37$ and $\Delta f_{\text{nl}}^{(\text{ort})} \sim 25$. More recently, by exploiting publicly available all-sky simulations, Lacasa et al. (2012) have tried to characterize the configuration dependence and the frequency behaviour of the bispectra originated by ‘radio’ and ‘far-infrared’ extragalactic sources. Their results show that the bispectrum due to (extragalactic) far-infrared sources starts to dominate that of radio sources on large angular scales at 150 GHz and dominates on the whole multipole range at 350 GHz.

In this paper we characterize the contamination that unresolved extragalactic point sources introduce in the CMB bispectrum by using the best available predictions on number counts of ‘radio’ sources, i.e. active galactic nuclei (AGN) with emission spectra dominated by synchrotron radiation (in this frequency range), and of ‘far-infrared’ (far-IR) sources, i.e. low- and high-redshift galaxies in which the microwave emission is dominated by warm/cold dust in thermal equilibrium with the radiation field. We also discuss the effect of unresolved *polarized* extragalactic sources in ‘future’ CMB anisotropy experiments. In particular, we make specific predictions for different frequency channels foreseen for the future *Cosmic Origins Explorer (CORE)* satellite mission (The CORE Collaboration et al. 2011), which has been recently submitted to the European Space Agency (ESA) Horizon 2015–2025 Programme.³

³ The *CORE* mission has not been selected by ESA in the 2011 first call for proposals. However, a *CORE*-like space mission with technical characteristics similar to the present ones could be hopefully funded in the near future, given that its very high scientific output has been already acknowledged by the ESA Space Science Advisory Committee (SSAC). In fact, a *CORE*-like mission will be of fundamental scientific importance for detecting the imprint left by primordial gravitational waves on polarized CMB anisotropies, since the amplitude of primordial gravitational waves is directly propor-

Our current predictions of number counts of extragalactic sources are based on two different models: the first one, by Tucci et al. (2011), is able to give the most precise predictions on high-frequency ($\nu \geq 100$ GHz) number counts of radio sources using physically motivated recipes to describe the complex spectral behaviour of blazars, that dominate the mm-wave counts at bright flux densities; the second one, by Lapi et al. (2011), using a simple physical model and a single reference spectral energy distribution (SED), has proved able to reproduce very well all the most recent measures [by *Herschel*, Balloon-borne Large Aperture Submillimeter Telescope (BLAST) etc.] of number counts of extragalactic sources at submm wavelengths. It is important to mention that, to our knowledge, this is the first time that predictions on the impact of extragalactic radio and far-infrared (IR)/submm sources on the CMB angular power spectrum, bispectrum and other statistics are performed with the level of detail discussed here.

In Section 2 we describe the extragalactic point source population in this frequency range. The models that characterize their power spectrum are presented in Section 3. Section 4 reviews the basic formalism for the angular bispectrum and the four bispectrum shapes. Section 5 introduces the point source bispectra and the estimator to measure the Δf_{nl} bias. The forecasts for the *CORE* mission are given in Section 6 and, finally, in Section 7 the conclusions are listed.

2 EXTRAGALACTIC POINT SOURCES

In this section we describe the expected contribution of extragalactic point sources to temperature and polarization fluctuations from centimetre to submillimetre wavelengths. In this range of frequencies, extragalactic sources are typically divided in two populations, according to the physical radiative mechanism involved: far-IR sources dominated by thermal dust emission and radio sources dominated by synchrotron emission. In the following analysis we use number counts predictions obtained by two very recent models: i.e. Tucci et al. (2011, hereafter T11), as for radio sources at cm to mm wavelengths and Lapi et al. (2011, hereafter L11), as for far-IR/submm sources. In both cases, these models provide predictions on source number counts and related statistics which have shown a very good agreement with observational data.

2.1 Radio sources

Recent evolutionary models for extragalactic radio sources (ERS) by de Zotti et al. (2005) and Massardi et al. (2010) are capable of providing a good fit to number counts from ~ 100 MHz and up to ~ 100 GHz. They adopt a simple power law, i.e. $S(\nu) \propto \nu^\alpha$, with an almost flat spectral index ($\alpha \simeq -0.1$) for extrapolating spectra of blazar sources⁴ at high frequencies. These models, however, overpredict the number counts of extragalactic synchrotron sources detected by the Atacama Cosmology Telescope (ACT) at 148 GHz (Marriage et al. 2011) and by *Planck* in all the High Frequency Instrument (HFI) channels (Planck Collaboration XIII 2011; Planck

tional to the energy scale at which the inflation occurred. We do not know the exact instrumental characteristics of this future space mission. However, we do useful predictions by using the current instrumental technical details quoted in the proposal (The CORE Collaboration et al. 2011).

⁴ Blazars are jet-dominated extragalactic objects, observed within a small angle of the jet axis and characterized by a highly variable non-thermal synchrotron emission at GHz frequencies, in which the beamed component dominates the observed emission (Angel & Stockman 1980).

Collaboration VII 2013). The main reason for this disagreement is the spectral steepening observed in *Planck* Early Release Compact Source Catalogue (ERCSC) sources at above ~ 70 GHz (Planck Collaboration VII 2011; Planck Collaboration XV 2011) and already suggested by other data sets (González-Nuevo et al. 2008; Sadler et al. 2008).

The more recent models by T11 use physically grounded recipes to describe the complex spectral behaviour of blazars, which dominate the mm-wavelengths number counts at bright flux densities. The main novelty of these models is the statistical prediction of a break frequency, ν_M , in spectra of blazars in agreement with classical physical models for the synchrotron emission produced in blazar jets (e.g., Blandford & Königl 1979). The most successful of these models is referred to as C2Ex.⁵ This model is able to give a very good fit to all the data of bright extragalactic radio sources available so far: number counts up to ~ 500 – 600 GHz and on spectral index distributions up to, at least, 200–300 GHz (see T11 and López-Caniego et al. 2012; Planck Collaboration VII 2013).

Polarization in ERS is typically observed to be a few per cent of total intensity at cm or mm wavelengths (e.g., Murphy et al. 2010; Battye et al. 2011; Sajina et al. 2011), with only very few objects showing a fractional polarization, $\Pi = P/S$, as high as ~ 10 per cent. We refer to the recent paper by Tucci & Toffolatti (2012) in order to characterize polarization properties of ERS and their contribution to power spectra. In that work authors show that the distribution of the fractional polarization for blazars is well described by a log-normal distribution with values of the median fractional polarization, Π_{med} , and of $\langle \Pi^2 \rangle^{1/2}$ depending on the frequency. At frequencies $\nu \gtrsim 40$ GHz we take $\Pi_{\text{med}} = 0.04, 0.03, 0.036$ and $\langle \Pi^2 \rangle^{1/2} = 0.06, 0.038, 0.045$ for steep-spectrum sources, FSRQs and BL Lacs, respectively. These values are intermediate between the two ‘extreme’ cases discussed in Tucci & Toffolatti (2012) and rely on the most recent data at 86 GHz of Agudo et al. (2010).

2.2 Far-IR sources

Recent data acquired by the *Herschel*-Astrophysical Terahertz Large Area Survey (H-ATLAS; Eales et al. 2010), the largest area survey carried out by the *Herschel Space Observatory* (Pilbratt et al. 2010), have demonstrated that far-IR sources comprise both a low- z galaxy population, identified through matching to the Sloan Digital Sky Survey (York et al. 2000) data (Smith et al. 2011), and a high- z population (median redshift ~ 2) identified through their far-IR colours (Amblard et al. 2010). Low- z galaxies are generally normal/star-forming late-type galaxies with moderate opacity (Dunne et al. 2011; Smith et al. 2012). Through analyses of clustering these two populations are found to be very different; the low- z population ($z < 0.3$) does cluster like star-forming blue galaxies (Maddox et al. 2010; Guo et al. 2011; van Kampen et al. 2012), whereas the high- z population clusters much more strongly, suggesting that the high- z sources reside in more massive haloes (Maddox et al. 2010).

In the case of late-type infrared galaxies, we adopt the same source number counts as in L11, previously derived by Negrello et al. (2007). As explained in more detail in that paper, the contributions of the this type of galaxies to the number counts were

estimated following a phenomenological approach, which consists of simple analytic recipes to evolve backwards their local luminosity functions (LFs) with cosmic time, as well as of appropriate templates for their SEDs, to compute K -corrections.

Regarding star-forming protospheroidal galaxies, the 100 and 250 μm LFs at different redshifts are quite well reproduced by the physical model of early-type galaxy (ETG) formation and evolution by Granato et al. (2001, 2004), further elaborated by Lapi et al. (2006), and recently revised by L11 without any adjustment of the parameters. As discussed in these papers, the model is built in the framework of the standard hierarchical clustering scenario. This model was the basis for the successful predictions of the submillimetre counts of strongly lensed galaxies by Negrello et al. (2007). It also accurately reproduced the epoch-dependent galaxy LFs in different spectral bands, as well as a variety of relationships among photometric, dynamical and chemical properties, as shown in previous papers (see table 2 of Lapi et al. 2006) and additional results, especially on the galaxy chemical evolution, in Mao et al. (2007) and Lapi et al. (2008).

Number counts from the L11 model are observed to be slightly high at mm wavelengths.⁶ The model, in fact, tends to overestimate the Poisson contribution to the power spectrum of intensity fluctuations measured at frequencies ≤ 353 GHz by *Planck* (Planck Collaboration XVIII 2011), South Pole Telescope (SPT; Hall et al. 2010) and ACT (Dunkley et al. 2011). Scaling down counts by constant factors is however enough to comply with measurements at mm wavelengths (Xia et al. 2012). In this work we multiply number counts by the following factors: 0.81 and 0.71 at 375 and 225 GHz, 0.55 for $\nu \leq 165$ GHz.

Finally, as predicted by Blain (1996), Perrotta et al. (2002, 2003), Negrello et al. (2007), Paciga, Scott & Chapin (2009) and Lima, Jain & Devlin (2010) among others, the combination of the very steep number counts in the far-IR/submm and of high redshifts – hence large lensing optical depths – maximizes the fraction of strongly lensed sources. In the case of the L11 model, the contribution from lensed protospheroidal galaxies were estimated using the amplification distribution of Perrotta et al. (2003), Negrello et al. (2007) and Lapi et al. (2012). Such predictions were recently confirmed by Negrello et al. (2010) and González-Nuevo et al. (2012) by using the Science Demonstration Phase data of the H-ATLAS survey. For the current work we have considered that the strong gravitational lensing effect is a random process, and therefore, the lensed protospheroidal galaxies show a Poisson spatial distribution. Because of their lack of correlation with the unlensed protospheroidal galaxy sky distribution, we have treated them as a different population.

Regarding polarization, all the – very few – data so far collected on the polarization of far-IR/submm sources (dusty early- or late-type galaxies and Galactic dust clouds) do indicate that the total polarized emission of these sources has to be very low with median values at around 1–2 per cent level. Moreover, in many observations, no polarized emission (or compatible with zero) has been detected (Dotson et al. 2010; Vaillancourt 2011). Because of the above, we have assumed – for the population of far-IR selected sources as a whole – a (conservative) average polarization level of 1 per cent.

⁵ It assumes different distributions of break frequencies for BL Lac objects and flat spectrum radio quasars (FSRQs) – with the relevant synchrotron emission coming from more compact regions in jets of the former objects. See T11 for more details.

⁶ This effect is due to the fact that the SED of high-redshift objects (that contribute to the mm counts around 10–30 mJy) are slightly different to the template adopted in L11 as a universal SED for protospheroids.

3 CONTAMINATION OF POINT SOURCES TO CMB POWER SPECTRA

In Fig. 1 we directly compare expected integrated number counts of radio and far-IR sources from the T11 and L11 models at frequencies ranging from 45 to 375 GHz – i.e. the frequencies interesting for CMB studies. The figure well emphasizes the different behaviour of the two populations in terms of flux density distribution and frequency spectrum. In the case of ERS, we can notice that the slope of integral counts ($N(>S) \propto S^{-\beta}$) is not very different from $\beta \simeq 1$ at $S < 1$ Jy. This means that the contamination of unresolved ERS to the CMB power spectra is almost proportional to the flux density limit, S_c , above which sources can be detected and removed or masked. Moreover, because the spectra of blazars are typically flat or moderately steep, $-1 \lesssim \alpha \lesssim 0$, integral counts change slowly with the frequency, decreasing by only a factor of 2–4 from 44 to 350 GHz, depending on the flux density considered.

On the other hand, far-IR source counts increase very steeply when moving to low flux densities, due to the contribution of protospheroidal galaxies (see dotted lines in Fig. 1). The general slope of far-IR integral counts, β , is between 2 and 3 depending on the flux density range. As a consequence, the amplitude of power spectra for unresolved far-IR sources is almost insensitive to the value of S_c , at least for flux density limits typically obtained in large-area CMB experiments. Finally, because of the steep frequency spectrum of dust emission, $N(>S)$ is also strongly dependent on frequency, making the contribution of far-IR sources already completely negligible at frequencies $\nu \lesssim 100$ GHz (see Fig. 2).

In terms of power spectra, it is well known that point sources without spatial correlations have

$$C_\ell^{\text{white}} = k_\nu^2 \int_0^{S_c} S^2 \frac{dn}{dS} dS, \quad (1)$$

where $k_\nu^{-1} = dB_0/dT_0$ is the conversion factor from flux density to temperature units (see e.g. Tegmark & Efstathiou 1996), S_c is the flux density limit for the unresolved point sources and dn/dS is number counts that characterize the population of sources. However, powerful high- z far-IR sources have been shown to be highly correlated by, e.g., recent measures from *Herschel* (Amblard et al.

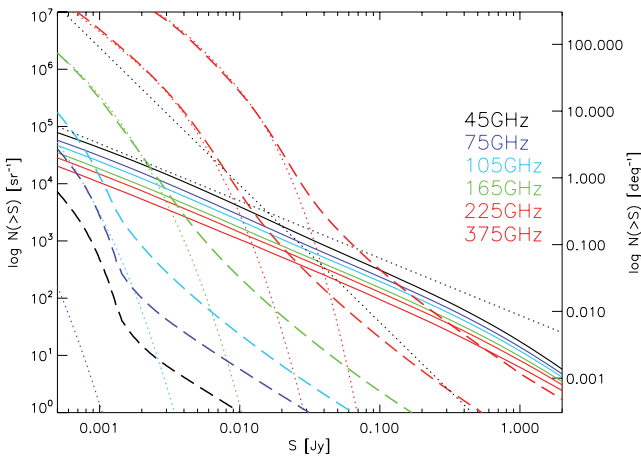


Figure 1. Predicted integral number counts of radio sources (solid lines; T11 model) and of far-IR sources (total counts: dashed lines; protospheroidal galaxies only: dotted lines; L11 model) at the frequencies indicated inside the panel (each frequency and the corresponding curves are displayed by a different colour). The two black dotted lines, plotted as a reference, represent two power laws, with slopes 1 and 3, respectively.

2011; Viero et al. 2012) and *Planck* (Planck Collaboration XVIII 2011). For the other extragalactic source populations the clustering is negligible (Toffolatti et al. 2005; Negrello et al. 2007). Therefore, we shall consider the contribution of clustered sources to angular power spectra (and bispectra) only for dusty protospheroidal galaxies. This is estimated from realistic simulations of two-dimensional distributions of clustered point sources (see González-Nuevo et al. 2005). This software was also implemented in the *Planck* Sky Model (Delabrouille et al. 2012) and its predictions on the power spectrum of cosmic IR background anisotropies have been shown to be in good agreement with the latest *Planck* measurements (Planck Collaboration XVIII 2011).

The total power spectra from point sources are then given by the sum of the following components:

$$C_\ell^{\text{ps}} = C_\ell^{\text{rad,white}} + C_\ell^{\text{ir,white}} + C_\ell^{\text{sph,clust}}, \quad (2)$$

where $C_\ell^{\text{rad,white}}$ and $C_\ell^{\text{ir,white}}$ are the constant term from shot-noise power spectra of radio and far-IR (which includes late-type, lensed and protospheroidal galaxies) sources, and $C_\ell^{\text{sph,clust}}$ is the power spectrum from the clustered contribution of protospheroidal galaxies.⁷

Regarding polarization, point sources contribute, on average, equally to power spectra of E- and B-modes or equivalently to Q and U power spectra in the flat-sky limit. Polarization power spectra of point sources are proportional to the temperature power spectrum (Tucci et al. 2004): for the Q power spectrum, e.g., we have

$$C_\ell^Q = k_\nu^2 \langle S^2 p^2 \cos^2(2\phi) \rangle = \frac{k_\nu^2}{2} \langle p^2 \rangle C_\ell^T. \quad (3)$$

The above relation assumes that the flux density S , the polarization fraction p and the polarization angle in the chosen reference system ϕ are independent variables (see Tucci & Toffolatti 2012, for a detailed discussion on this assumption). On the other hand, all the cross-power spectra are equal to zero:

$$C_\ell^{TQ} \propto \langle S^2 p \cos(2\phi) \rangle = 0,$$

$$C_\ell^{TU} \propto \langle S^2 p \sin(2\phi) \rangle = 0,$$

$$C_\ell^{QU} \propto \langle S^2 p^2 \cos(2\phi) \sin(2\phi) \rangle = 0. \quad (4)$$

Fig. 2 shows model predictions for temperature and polarization power spectra at 75, 105, 165 and 225 GHz. We also consider two possible flux density limits, i.e. $S_c = 1$ and 0.3 Jy.

As previously discussed, far-IR spectra are almost insensitive to the choice of S_c , especially for the clustered contribution. Comparing the two source populations, we see that radio sources are dominant at $\nu \lesssim 100$ GHz, while far-IR sources start to dominate at $\nu \gtrsim 200$ GHz both in temperature and polarization. At intermediate frequencies, i.e. $100 < \nu < 200$ GHz, ERS are the main contaminant at high multipoles ($\ell > 1000$), but not for $\ell \lesssim 1000$, where the far-IR clustered sources contribution becomes higher than the ERS power spectrum in temperature and comparable in polarization.

Because protospheroidal galaxies are the only clustered population in the L11 model, it is worth discussing more in detail their relevance in terms of power spectrum. In Fig. 3 we plot the ratio of the total spectrum of protospheroidal galaxies to their shot-noise spectrum, $C_\ell^{\text{sph}}/C_\ell^{\text{sph,white}}$, and to the shot-noise spectrum of all far-IR sources, $C_\ell^{\text{sph}}/C_\ell^{\text{ir,white}}$ (the former ratio enters later in equation 18

⁷ We remind the reader that the total power spectrum of a population of clustered sources is $C_\ell^{\text{tot}} = C_\ell^{\text{white}} + C_\ell^{\text{clust}}$.

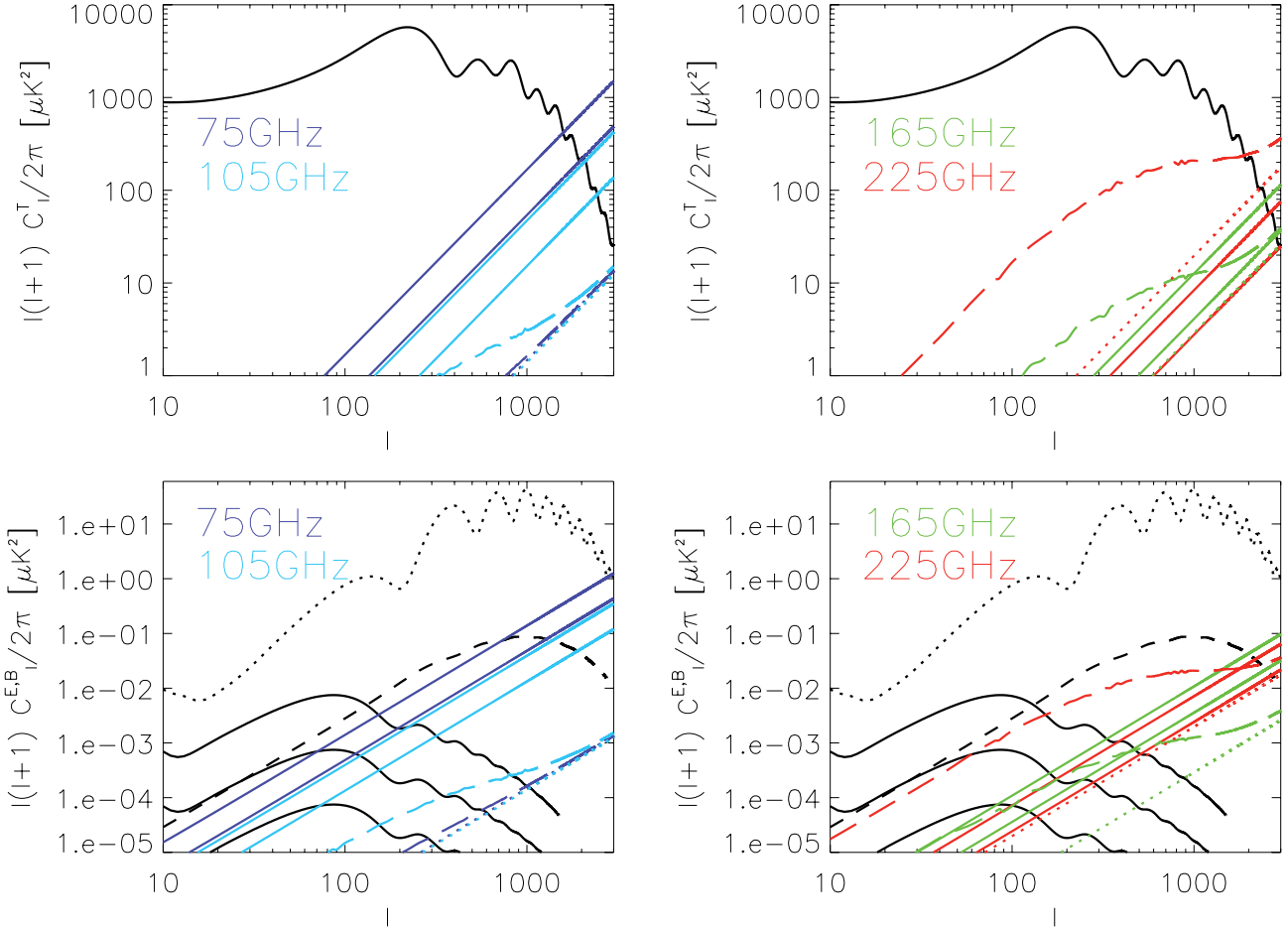


Figure 2. Point source power spectra for temperature (top panels) and polarization (bottom panels) at 75, 105, 165 and 225 GHz. Lines in colour are for radio sources (solid lines) and for far-IR sources (dashed lines for the total contribution; dotted lines for the Poisson contribution). We consider two flux density limits for radio sources: $S_c = 1$ and 0.3 Jy; for far-IR sources we plot only the case with $S_c = 1$ (there are not significant differences taking $S_c = 0.3$). Black lines are for CMB spectra from the fiducial *WMAP* 7-year model: in the bottom panels dotted lines are for the E-mode, dashed lines for lensing-induced B-mode and solid lines for primordial B-mode with $r = 0.1, 0.01, 10^{-3}$.

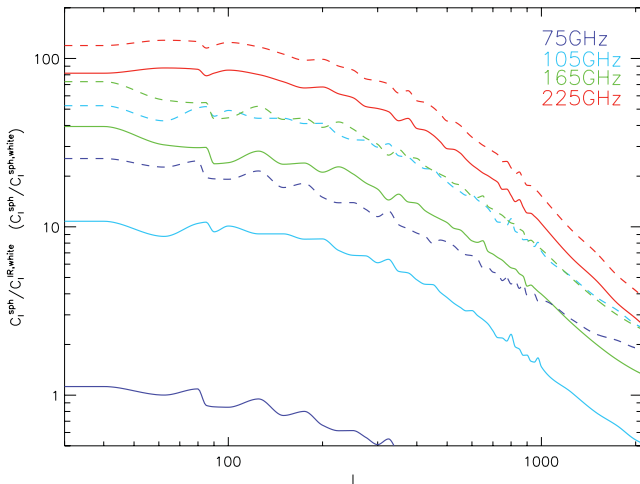


Figure 3. The ratio of the total spectrum, C_l^{sph} , to the shot-noise spectrum, $C_l^{\text{sph,white}}$, of protospheroidal galaxies (dashed lines), and to the shot-noise spectrum from far-IR sources, $C_l^{\text{ir,white}}$ (solid lines), at the frequencies indicated in the panel.

and gives us the increase of the bispectrum due to clustering). As the frequency decreases, we observe two effects: (a) the contribution of protospheroidal galaxies to the far-IR power spectrum becomes less and less important; (b) moreover, the relative contribution of clustering to protospheroidal power spectrum reduces. The reason of these trends is related to the shift of the redshift distribution of submillimetre bright galaxies to higher redshift with increasing the wavelength, as effect of the strongly negative K -correction (L11). At high z , however, massive protospheroidal galaxies become increasingly rare and the slope of number counts at low frequencies is not as steep as that at submm wavelengths. As a consequence, the shot-noise term of the power spectrum, dominated by bright sources, has more weight with respect to the clustered one (dominated by faint sources).

Finally, we discuss the contamination of extragalactic sources to CMB fluctuations: in the frequency range considered, it is relevant only at very high l ($\gtrsim 2000$) for temperature and is essentially negligible for the E-mode polarization (see Fig. 2). This is not the case for the B-mode polarization arisen from primordial gravitational waves if $r < 0.01$. A subtraction of points sources at levels even lower than ones shown in Fig. 2 is not an easy task. For radio sources, this could be in principle done (a) by removing/masking sources down to flux density limits of $S < 0.3$ Jy or (b) by a direct measurement

of the power spectrum of polarization fluctuations at small angular scales. In both cases observations with arcmin resolution are required. Far-IR sources are still more challenging to deal with, and the previous methods are not valid to clean their contamination due to the high correlated signal. Instead, approaches similar to ones used for diffuse Galactic foregrounds should be more suitable for this goal, although they are model dependent.

4 THE ANGULAR BISPECTRUM

The CMB angular bispectrum is defined in terms of cubic combinations of the spherical harmonic coefficients of the CMB anisotropies of the T , E and B modes:

$$B_{\ell_1 \ell_2 \ell_3, m_1 m_2 m_3}^{ijk} \equiv \left\langle a_{\ell_1 m_1}^i a_{\ell_2 m_2}^j a_{\ell_3 m_3}^k \right\rangle, \quad (5)$$

where the indices $\{i, j, k\}$ cover the T, E and B modes. The angular-averaged bispectrum and the reduced angular-averaged bispectrum, $B_{\ell_1 \ell_2 \ell_3}^{ijk}$ and $b_{\ell_1 \ell_2 \ell_3}^{ijk}$ are defined as

$$\begin{aligned} B_{\ell_1 \ell_2 \ell_3}^{ijk} &= \sum_{m_1 m_2 m_3} \begin{pmatrix} \ell_1 & \ell_2 & \ell_3 \\ m_1 & m_2 & m_3 \end{pmatrix} B_{\ell_1 \ell_2 \ell_3, m_1 m_2 m_3}^{ijk} \\ &= b_{\ell_1 \ell_2 \ell_3}^{ijk} \begin{pmatrix} \ell_1 & \ell_2 & \ell_3 \\ 0 & 0 & 0 \end{pmatrix} \sqrt{\frac{(2\ell_1 + 1)(2\ell_2 + 1)(2\ell_3 + 1)}{4\pi}}. \end{aligned} \quad (6)$$

The reduced bispectrum can be written as (see e.g. Yadav & Wandelt 2010)

$$\begin{aligned} b_{\ell_1 \ell_2 \ell_3}^{ijk} &= (4\pi)^3 (-i)^{\ell_1 + \ell_2 + \ell_3} \sum_{m_1 m_2 m_3} \begin{pmatrix} \ell_1 & \ell_2 & \ell_3 \\ m_1 & m_2 & m_3 \end{pmatrix} \\ &\times \int \frac{d^3 k_1}{(2\pi)^3} \frac{d^3 k_2}{(2\pi)^3} \frac{d^3 k_3}{(2\pi)^3} Y_{\ell_1 m_1}^*(\hat{k}_1) Y_{\ell_2 m_2}^*(\hat{k}_2) Y_{\ell_3 m_3}^*(\hat{k}_3) \\ &\times g_{\ell_1}^i(k_1) g_{\ell_2}^j(k_2) g_{\ell_3}^k(k_3) \langle \Phi(\mathbf{k}_1) \Phi(\mathbf{k}_2) \Phi(\mathbf{k}_3) \rangle, \end{aligned} \quad (7)$$

where $g_\ell^i(k)$ is the transfer function for temperature or polarization⁸ and $\langle \Phi(\mathbf{k}_1) \Phi(\mathbf{k}_2) \Phi(\mathbf{k}_3) \rangle$ is the three-point correlation function in Fourier space of the Bardeen's curvature perturbations $\Phi(\mathbf{x})$. The three-point correlation function can be given in terms of the so-called shape function $F(k_1, k_2, k_3)$:

$$\langle \Phi(\mathbf{k}_1) \Phi(\mathbf{k}_2) \Phi(\mathbf{k}_3) \rangle = (2\pi)^3 \delta^3(\mathbf{k}_1 + \mathbf{k}_2 + \mathbf{k}_3) F(k_1, k_2, k_3). \quad (8)$$

The level of primordial non-Gaussianity is parametrized by the nonlinear coupling parameter f_{nl} (Verde et al. 2000; Komatsu & Spergel 2001; Bartolo et al. 2004; Yadav & Wandelt 2010). This parameter measures departures from zero of the bispectrum. Depending on the physical mechanisms of the different inflationary models the shape function can take different forms. The shapes that we consider in this work are the local, equilateral, orthogonal and flat shapes.

Local shape. This type of Non-Gaussianity is generated in multifield inflationary models (Komatsu, Spergel & Wandelt 2005; Komatsu 2010), the curvaton (Lyth, Ungarelli & Wands 2003), the inhomogeneous reheating scenario (Bartolo et al. 2004; Dvali et al.

2004), hybrid inflation (Lin 2009) etc. Its shape is given by (see e.g. Creminelli et al. 2006; Yadav & Wandelt 2010)

$$F(k_1, k_2, k_3) = 2A^2 f_{\text{nl}} \left[\frac{1}{k_1^{3-(n_s-1)} k_2^{3-(n_s-1)}} + (2 \text{ perm}) \right], \quad (9)$$

where A is the amplitude of the power spectrum and n_s is the spectral index (Komatsu et al. 2011). Most of the signal of this type of bispectrum is located in squeezed configurations such as $k_1 \simeq k_2 \gg k_3$ (Creminelli et al. 2006; Yadav & Wandelt 2010; Martínez-González & Planck Collaboration 2012).

Equilateral shape. This type of non-Gaussianity is generated in the Dirac-Born-Infeld inflation (Bartolo et al. 2004; Silverstein & Tong 2004; Langlois et al. 2008), ghost inflation (Arkani-Hamed et al. 2004), single-field inflationary models in Einstein gravity (Chen et al. 2007) etc. Its shape is given by (see e.g. Creminelli et al. 2006; Yadav & Wandelt 2010)

$$\begin{aligned} F(k_1, k_2, k_3) &= 6A^2 f_{\text{nl}} \left[-\frac{1}{k_1^{3-(n_s-1)} k_2^{3-(n_s-1)}} + (2 \text{ perm}) \right. \\ &\quad - \frac{2}{(k_1 k_2 k_3)^{2(4-n_s)/3}} \\ &\quad \left. + \left\{ \frac{1}{k_1^{(4-n_s)/3} k_2^{2(4-n_s)/3} k_3^{(4-n_s)}} + (5 \text{ perm}) \right\} \right]. \end{aligned} \quad (10)$$

Most of the signal of this type of bispectrum is located in equilateral configurations such as $k_1 \simeq k_2 \simeq k_3$ (see e.g. Creminelli et al. 2006; Yadav & Wandelt 2010; Martínez-González & Planck Collaboration 2012).

Orthogonal shape. This type of Non-Gaussianity is generated in general single-field models (Cheung et al. 2008; Senatore, Smith & Zaldarriaga 2010). Its shape is given by (see e.g. Senatore et al. 2010; Yadav & Wandelt 2010; Komatsu et al. 2011)

$$\begin{aligned} F(k_1, k_2, k_3) &= 6A^2 f_{\text{nl}} \left[-\frac{3}{k_1^{3-(n_s-1)} k_2^{3-(n_s-1)}} + (2 \text{ perm}) \right. \\ &\quad - \frac{8}{(k_1 k_2 k_3)^{2(4-n_s)/3}} \\ &\quad \left. + \left\{ \frac{3}{k_1^{(4-n_s)/3} k_2^{2(4-n_s)/3} k_3^{(4-n_s)}} + (5 \text{ perm}) \right\} \right]. \end{aligned} \quad (11)$$

This bispectrum peaks in both equilateral $k_1 \simeq k_2 \simeq k_3$ and flat-triangle configurations such as $k_1 = k_2 + k_3$ (see the shape figures e.g. in Senatore et al. 2010; Martínez-González & Planck Collaboration 2012).

Flat shape. This type of non-Gaussianity is generated for example in models with modifications in the initial state of the inflaton field (see e.g. Meerburg, van der Schaar & Stefano Corasaniti 2009, and references therein), single-field models with high derivative iterations (Bartolo et al. 2010b) among others. Its shape is given by (Meerburg et al. 2009)

$$\begin{aligned} F(k_1, k_2, k_3) &= 6A^2 f_{\text{nl}} \left[\frac{1}{k_1^{3-(n_s-1)} k_2^{3-(n_s-1)}} + (2 \text{ perm}) \right. \\ &\quad + \frac{3}{(k_1 k_2 k_3)^{2(4-n_s)}} \\ &\quad \left. - \left\{ \frac{1}{k_1^{(4-n_s)/3} k_2^{2(4-n_s)/3} k_3^{(4-n_s)}} + (5 \text{ perm}) \right\} \right]. \end{aligned} \quad (12)$$

This bispectrum peaks in flat-triangle configurations $k_1 = k_2 + k_3$ (see the shape figure in Meerburg et al. 2009).

⁸ This function can be computed numerically using the CAMB (Lewis et al. 2000) software <http://camb.info>

Using the shape forms defined in equations (9)–(12) and the CMB bispectrum in equation (7) we have

$$(b_{\ell_1 \ell_2 \ell_3}^{\text{loc}})^{ijk} = 2 \int_0^\infty x^2 dx \left[\alpha_{\ell_1}^i(x) \beta_{\ell_2}^j(x) \beta_{\ell_3}^k(x) + (2 \text{ perm}) \right], \quad (13)$$

$$(b_{\ell_1 \ell_2 \ell_3}^{\text{eq}})^{ijk} = 6 \int_0^\infty dx x^2 \left[-\alpha_{\ell_1}^i(x) \beta_{\ell_2}^j(x) \beta_{\ell_3}^k(x) + (2 \text{ perm}) \right. \\ \left. + \beta_{\ell_1}^i(x) \gamma_{\ell_2}^j(x) \delta_{\ell_3}^k(x) + (5 \text{ perm}) - 2 \delta_{\ell_1}^i(x) \delta_{\ell_2}^j(x) \delta_{\ell_3}^k(x) \right], \quad (14)$$

$$(b_{\ell_1 \ell_2 \ell_3}^{\text{ort}})^{ijk} = 18 \int_0^\infty dx x^2 \left[-\alpha_{\ell_1}^i(x) \beta_{\ell_2}^j(x) \beta_{\ell_3}^k(x) + (2 \text{ perm}) \right. \\ \left. + \beta_{\ell_1}^i(x) \gamma_{\ell_2}^j(x) \delta_{\ell_3}^k(x) + (5 \text{ perm}) - \frac{8}{3} \delta_{\ell_1}^i(x) \delta_{\ell_2}^j(x) \delta_{\ell_3}^k(x) \right] \quad (15)$$

and

$$(b_{\ell_1 \ell_2 \ell_3}^{\text{flat}})^{ijk} = 6 \int_0^\infty dx x^2 \left[\alpha_{\ell_1}^i(x) \beta_{\ell_2}^j(x) \beta_{\ell_3}^k(x) + (2 \text{ perm}) \right. \\ \left. - \beta_{\ell_1}^i(x) \gamma_{\ell_2}^j(x) \delta_{\ell_3}^k(x) + (5 \text{ perm}) + 3 \delta_{\ell_1}^i(x) \delta_{\ell_2}^j(x) \delta_{\ell_3}^k(x) \right], \quad (16)$$

where $\alpha_\ell^i(x)$, $\beta_\ell^i(x)$, $\gamma_\ell^i(x)$, $\delta_\ell^i(x)$ are filter functions defined in terms of the power spectrum $P(k)$ and the transfer functions $g_\ell^i(k)$ (see e.g. Fergusson, Liguori & Shellard 2010a,b; Yadav & Wandelt 2010; Curto et al. 2011, for their actual definitions).

5 BIAS AND UNCERTAINTIES ON PRIMORDIAL NON-GAUSSIANITY

5.1 Bispectrum of point sources

The bispectrum of a point source population without spatial correlations, as the case of radio sources, is well characterized and is constant with ℓ :

$$b_{\ell_1 \ell_2 \ell_3}^{\text{white}} = k_v^3 \int_0^{S_c} S^3 \frac{dn}{dS} dS. \quad (17)$$

On the other hand, far-IR sources have spatial correlations. According to the widely used prescription by Argüeso et al. (2003), also

considered e.g. in González-Nuevo et al. (2005) and Argüeso et al. (2006) and extended to the full sky by Lacasa et al. (2012), the total bispectrum for a single clustered sources population can be written as

$$b_{\ell_1 \ell_2 \ell_3}^{\text{tot}} \equiv b_{\ell_1 \ell_2 \ell_3}^{\text{white}} \sqrt{\frac{C_{\ell_1}^{\text{tot}} C_{\ell_2}^{\text{tot}} C_{\ell_3}^{\text{tot}}}{C_{\ell_1}^{\text{white}} C_{\ell_2}^{\text{white}} C_{\ell_3}^{\text{white}}}}. \quad (18)$$

The total bispectrum of point sources is then given by the sum of the shot-noise contributions of ERS ($b_{\ell_1 \ell_2 \ell_3}^{\text{rad, white}}$) and of far-IR late-type and lensed galaxies ($b_{\ell_1 \ell_2 \ell_3}^{\text{lt+le, white}}$) plus the ‘clustered’ bispectrum of protospheroidal galaxies ($b_{\ell_1 \ell_2 \ell_3}^{\text{sph}}$) as given by equation (18):

$$b_{\ell_1 \ell_2 \ell_3}^{\text{ps}} = b_{\ell_1 \ell_2 \ell_3}^{\text{rad, white}} + b_{\ell_1 \ell_2 \ell_3}^{\text{lt+le, white}} + b_{\ell_1 \ell_2 \ell_3}^{\text{sph}}. \quad (19)$$

When polarization is considered, we have 10 different bispectra arisen from all the possible combinations of temperature and polarization E - and B -modes. However, on the base of the discussion in Section 3, it is easy to find out that the cross-bispectrum $b_{\ell_1 \ell_2 \ell_3}^{\text{TEB}}$ and all bispectra with an odd number of polarization terms are zero, i.e. $b_{\ell_1 \ell_2 \ell_3}^{\text{EEE}} = b_{\ell_1 \ell_2 \ell_3}^{\text{TTE}} = \dots = 0$. For the remaining two cases, the bispectrum is proportional to $b_{\ell_1 \ell_2 \ell_3}^{\text{TTT}}$ and it is given as

$$b_{\ell_1 \ell_2 \ell_3}^{\text{TEE}} = b_{\ell_1 \ell_2 \ell_3}^{\text{TBB}} = \frac{1}{2} \langle p^2 \rangle b_{\ell_1 \ell_2 \ell_3}^{\text{TTT}}. \quad (20)$$

The total power spectrum and the bispectrum of the point sources are the two ingredients needed to provide an estimate of their impact on the primordial non-Gaussianity f_{nl} parameter. The shot-noise terms of TTT and TEE bispectra from radio and far-IR sources are shown in Fig. 4. The frequency at which far-IR sources equal or overcome the ERS bispectrum is shifted to higher frequencies with respect to the power spectrum, i.e. it is around 300 (350) GHz for the TTT (TEE) bispectrum. This is in agreement with the fact that sources with flux density close to S_c give a larger contribution to the bispectrum than to the power spectrum. A significant dependence on the flux density limit is observed in the bispectrum of ERS and of late-type plus lensed galaxies, but not of protospheroidal galaxies.

Finally, we compare the total bispectrum of far-IR sources with the ERS one as a function of multipoles and for the four used configurations (see Fig. 5). The ratio is $\gtrsim 1$ only at 225 GHz for TTT , and for TEE , but at low multipoles. Otherwise, far-IR sources are always subdominant, or even negligible, unless considering multipoles

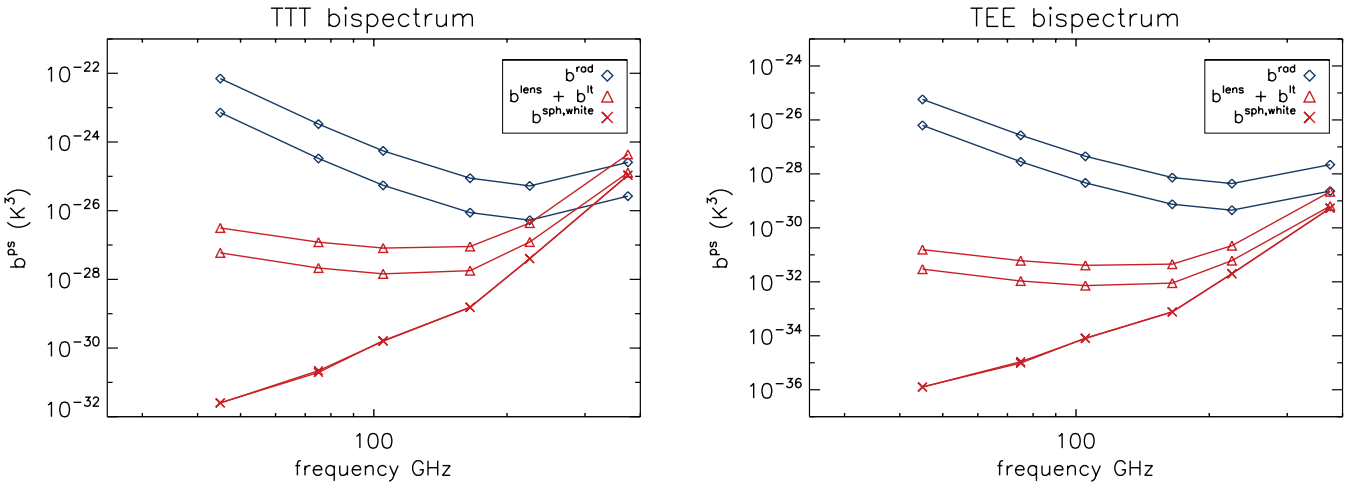


Figure 4. The amplitude of the shot-noise bispectrum of radio sources (b^{rad}) and of the far-IR source (for late-type plus lensed galaxies, $b^{\text{lt+le}}$, and for protospheroidal galaxies, $b^{\text{sph, white}}$) as a function of frequency and for the two considered flux density limits, i.e. $S_c = 0.3$ and 1 Jy. The left-hand panel corresponds to the TTT bispectrum and the right-hand panel corresponds to the TEE bispectrum.

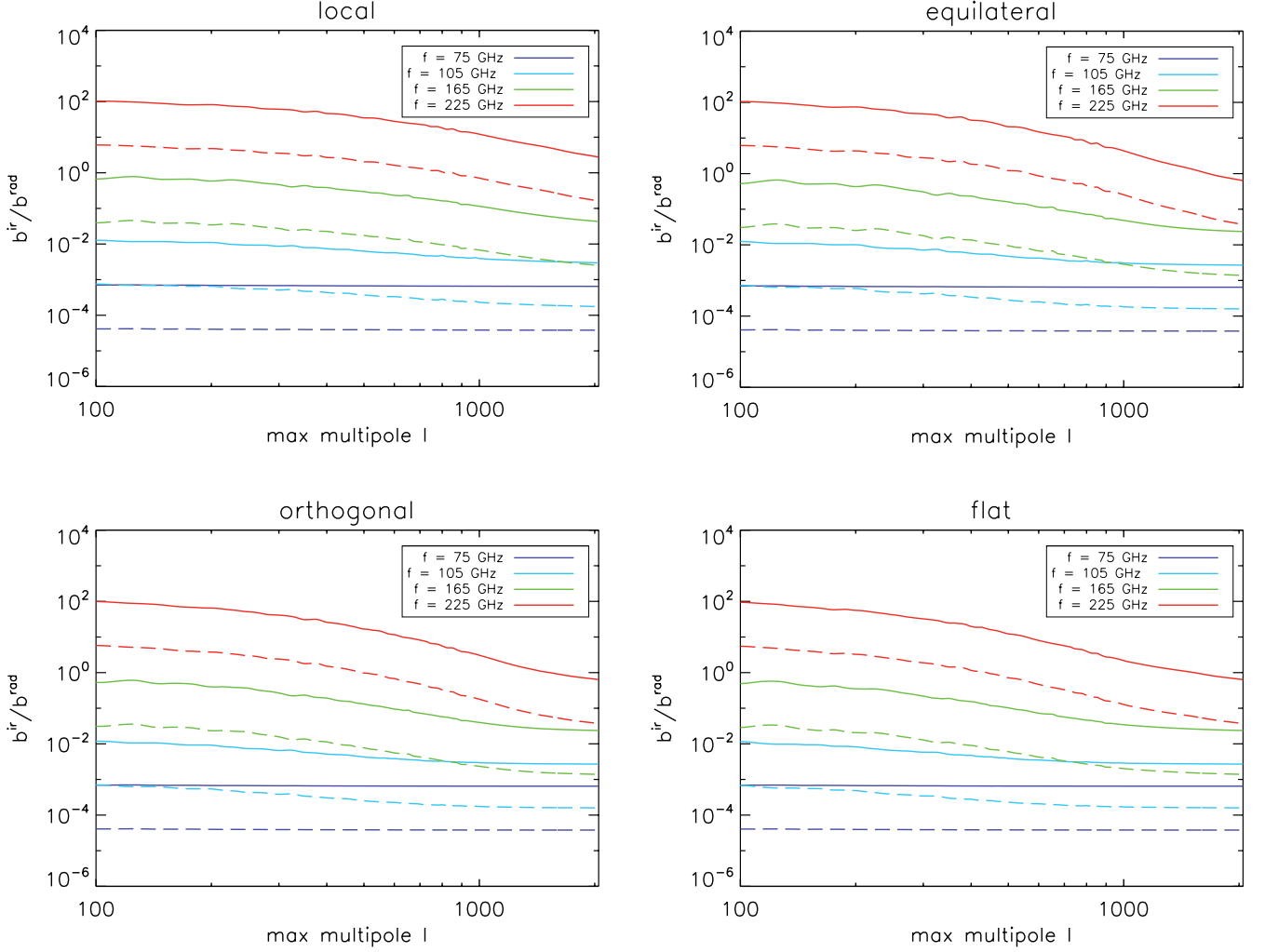


Figure 5. Ratio of the far-IR source bispectrum b^{ir} to the radio source bispectrum b^{rad} for $S_c = 0.3$ Jy and the four considered shapes: local configuration ($\ell_1 = \ell_2 = \ell$ and $\ell_3 = \ell_{\min} = 2$); equilateral configuration ($\ell_1 = \ell_2 = \ell_3 = \ell$); isosceles orthogonal configuration ($\ell_1 = \ell_2 = \ell$ and $\ell_3 = \sqrt{2}\ell$) and isosceles flat configuration ($\ell_1 = \ell_2 = \ell$ and $\ell_3 = 2\ell$). The solid lines correspond to T , and the long dashed lines correspond to T and E .

$\ell < 100$. As widely discussed in Section 3, clustering of far-IR sources becomes less relevant at low frequencies, and at $\nu \lesssim 100$ GHz the far-IR bispectrum is almost constant. Fig. 5 can be directly compared to fig. 9 of Lacasa et al. (2012). Various differences can be highlighted from the two plots: (1) at high multipoles, the bispectrum ratio we find is typically higher than one order of magnitude; (2) far-IR bispectra by Lacasa et al. (2012) have a steeper slope; (3) no frequency dependence is observed in Lacasa et al. (2012) results. These discrepancies should be attributed to the different way to model far-IR sources and their correlation.

5.2 Bias and uncertainties on f_{nl}

Considering weak levels of non-Gaussianity, the bias induced in the primordial bispectrum $B_{\ell_1 \ell_2 \ell_3}^{\text{prim}}$ by the point-source bispectrum $B_{\ell_1 \ell_2 \ell_3}^{\text{ps}}$ is (see e.g. Komatsu & Spergel 2001; Bartolo et al. 2004; Yadav & Wandelt 2010; Lacasa et al. 2012; Martínez-González & Planck Collaboration 2012)

$$\Delta f_{\text{nl}} = \sigma^2(f_{\text{nl}}) \sum_{\ell_1 \leq \ell_2 \leq \ell_3 \leq \ell_{\max}} \frac{B_{\ell_1 \ell_2 \ell_3}^{\text{ps}} B_{\ell_1 \ell_2 \ell_3}^{\text{prim}}}{\left(\Delta_{\ell_1 \ell_2 \ell_3} C_{\ell_1}^{\text{tot}} C_{\ell_2}^{\text{tot}} C_{\ell_3}^{\text{tot}} \right)}, \quad (21)$$

where

$$C_{\ell}^{\text{tot}} = \left(C_{\ell}^{\text{CMB}} + C_{\ell}^{\text{ps}} \right) w_{\ell}^2 + C_{\ell}^{\text{noise}}, \quad (22)$$

C_{ℓ}^{CMB} is the CMB power spectrum, w_{ℓ} contains the instrumental and pixelization beams, C_{ℓ}^{noise} is the power spectrum of the noise, $\Delta_{\ell_1 \ell_2 \ell_3} = 1 + 2\delta_{\ell_1 \ell_2} \delta_{\ell_2 \ell_3} + \delta_{\ell_1 \ell_2} + \delta_{\ell_2 \ell_3} + \delta_{\ell_1 \ell_3}$. The error-bar of the f_{nl} parameter is given by (Komatsu & Spergel 2001)

$$\sigma^{-2}(f_{\text{nl}}) = \sum_{\ell_1 \leq \ell_2 \leq \ell_3} \frac{\left(B_{\ell_1 \ell_2 \ell_3}^{\text{prim}} \right)^2}{\left(\Delta_{\ell_1 \ell_2 \ell_3} C_{\ell_1}^{\text{tot}} C_{\ell_2}^{\text{tot}} C_{\ell_3}^{\text{tot}} \right)}. \quad (23)$$

This estimator can be generalized to include the polarization:

$$\Delta f_{\text{nl}} = \sigma^2(f_{\text{nl}}) \times \sum_{ijkrst} \sum_{\ell_1 \leq \ell_2 \leq \ell_3 \leq \ell_{\max}} \left(B_{\ell_1 \ell_2 \ell_3}^{\text{ps}} \right)^{ijk} C_{ijk,rst}^{-1} \left(B_{\ell_1 \ell_2 \ell_3}^{\text{prim}} \right)^{rst}, \quad (24)$$

where

$$\sigma^{-2}(f_{\text{nl}}) = \sum_{ijkrst} \sum_{\ell_1 \leq \ell_2 \leq \ell_3 \leq \ell_{\max}} \left(B_{\ell_1 \ell_2 \ell_3}^{\text{prim}} \right)^{ijk} C_{ijk,rst}^{-1} \left(B_{\ell_1 \ell_2 \ell_3}^{\text{prim}} \right)^{rst}. \quad (25)$$

The covariance matrix $C_{ijk, rst}$ is given by (Yadav, Komatsu & Wandelt 2007)

$$\begin{aligned} C_{ijk, prq} = & C_{\ell_1}^{ip} C_{\ell_2}^{jq} C_{\ell_3}^{kr} + C_{\ell_1}^{ip} C_{\ell_2}^{jr} C_{\ell_3}^{kq} \delta_{\ell_2 \ell_3} + C_{\ell_1}^{iq} C_{\ell_2}^{jp} C_{\ell_3}^{kr} \delta_{\ell_1 \ell_2} \\ & + C_{\ell_1}^{ir} C_{\ell_2}^{jq} C_{\ell_3}^{kp} \delta_{\ell_1 \ell_3} + C_{\ell_1}^{iq} C_{\ell_2}^{jr} C_{\ell_3}^{kp} \delta_{\ell_1 \ell_2} \delta_{\ell_2 \ell_3} \delta_{\ell_3 \ell_1} \\ & + C_{\ell_1}^{ir} C_{\ell_2}^{jp} C_{\ell_3}^{kq} \delta_{\ell_1 \ell_3} \delta_{\ell_2 \ell_1} \delta_{\ell_2 \ell_3}. \end{aligned} \quad (26)$$

The indices $\{i, j, k, r, s, t\}$ cover the intensity T and the polarization E and B modes. The cosmological signal for the B mode is negligible. Including the point source B mode signal in the estimator does not improve the estimates of $\sigma(f_{\text{nl}})$ and $\Delta(f_{\text{nl}})$. Therefore we just include in the estimator the T and E modes for the case with polarization. As in the case with intensity only, each of the power or cross-power spectrum in equation (26) contains the CMB signal, instrumental beam, instrumental noise spectrum and the point source spectrum contributions.

6 RESULTS

We have studied the impact of unresolved point sources on the possible measurement of the f_{nl} parameter for two cases: (a) an ideal experiment without instrumental beam nor instrumental noise; (b) a possible future *CORe*-like mission, sensitive to temperature and polarization, with instrumental parameters adopted from The CORe Collaboration et al. (2011) and summarized in Table 1. The CMB power spectrum and transfer function for temperature and polarization needed to estimate the f_{nl} bias and its uncertainty (equations 21, 23–25) are numerically computed with CAMB (Lewis et al. 2000) using the cosmological parameters that best-fits *WMAP* 7-year data. In this section we present results for the uncertainty on the f_{nl} parameter, $\sigma(f_{\text{nl}})$, for the expected bias, $\Delta(f_{\text{nl}})$, and for the relative bias, $\Delta(f_{\text{nl}})/\sigma(f_{\text{nl}})$ (see Fig. 6, Tables 2 and 3 for the ideal case, and Fig. 7, Tables 4 and 5 for a *CORe*-like mission). We compare results obtained with temperature only and with temperature plus polarization. Four different f_{nl} shapes are taken into account, i.e. local, equilateral, orthogonal and flat. Figs 6 and 7 are for $S_c = 0.3$ Jy and for frequencies 75, 105, 165 and 225 GHz. In the tables we also provide the results, evaluated at $\ell_{\text{max}} = 2048$, for $S_c = 1.0$ Jy and for 45 and 375 GHz.

The following results assume full sky. This is an unrealistic hypothesis as Galactic plane and strong point sources have to be masked. How $\sigma(f_{\text{nl}})$ depends on the mask and on the fraction of the sky available, f_{sky} , is something complex and outside the purpose of this work. However, for large areas of the sky ($f_{\text{sky}} \gtrsim 0.7$) the following approximation can be used (Komatsu et al. 2002):

$$\sigma(f_{\text{nl}})_{\text{cutsky}} = \sigma(f_{\text{nl}})_{\text{fullsky}} / \sqrt{f_{\text{sky}}}. \quad (27)$$

The bias on f_{nl} does not depend on f_{sky} , as it can be seen from equation (21).

Table 1. Instrumental properties for a possible future *CORe*-like mission, based on values given by The CORe Collaboration et al. (2011).

ν (GHz)	45	75	105	165	225	375
FWHM (arcmin)	23.3	14.0	10.0	6.4	4.7	2.8
T rms noise ^a	5.25	2.73	2.68	2.67	2.64	68.6
Pol. rms noise ^a	9.07	4.72	4.63	4.61	4.57	119

^a μK arcmin.

6.1 Uncertainty

In equations (23) and (25) of the previous section we have provided the analytical formulas to estimate the uncertainty on the f_{nl} parameter. As we can see, the point sources contribution enters only in the total angular power spectrum (equation 22), in the sum $C_{\ell}^{\text{CMB}} + C_{\ell}^{\text{ps}}$. In the frequency range 75–225 GHz, C_{ℓ}^{ps} is always subdominant with respect to T and E -mode CMB spectra up to $\ell \sim 2000$ (see Fig. 2). Therefore, as illustrated by Figs 6 and 7, the impact of point sources on the uncertainty in the determination of the f_{nl} parameter is small or negligible, especially when polarization is considered. In this case, in fact, $\sigma(f_{\text{nl}})$ is almost indistinguishable with the intrinsic f_{nl} uncertainty due to the cosmic variance (black curves in the plots, indicated as ‘no sources’). This is true also for the ‘temperature only’ case at 105 and 165 GHz, i.e. the frequencies where the signal from point sources is minimum. At 75 GHz the uncertainty becomes slightly higher only at $\ell \gtrsim 1500$ because of the increasing contribution of ERS; while, a larger $\sigma(f_{\text{nl}})$ is obtained at 225 GHz in the all range of ℓ s, clearly due to the clustered signal of IR sources. At 353 GHz the point sources contamination grows dramatically, and the $\sigma(f_{\text{nl}})$ is around four to six times higher than at 225 GHz. On the other hand, the loss of sensitivity between 75 and 45 GHz is not so strong, due to the flatter spectral behaviour of ERS. These results seem to be quite independent of the two flux density limits considered, even at 45 and 75 GHz, where $\sigma(f_{\text{nl}})$ is about 10–20 per cent larger (see Tables 2–5).

When polarization is taken into account, f_{nl} error-bars are reduced approximately by a factor of 2–3 accordingly to the shape, in the frequency range 75–225 GHz. At 45 and 375 GHz, polarization permits to achieve a very large improvement in the f_{nl} error-bars, but only for the ideal case. In a *CORe*-like mission the instrumental noise in polarization saturates the f_{nl} error-bars at these frequencies and prevents to reduce them as in the ideal case.

The values of $\sigma(f_{\text{nl}})$ obtained from a *CORe*-like mission are very close to the results of the ideal case. This means that observational performances, in terms of instrumental noise and beam, given in Table 1, are almost optimal for f_{nl} measurements, at least up to $\ell \sim 2000$. Only in the 75-GHz channel (and partially in the 105-GHz channel) the effect of the beam is observed in Fig. 7, where $\sigma(f_{\text{nl}})$ starts to be constant at $\ell_{\text{max}} \gtrsim 1500$. In the 165- and 225-GHz channels the f_{nl} uncertainty is always decreasing up to $\ell_{\text{max}} = 2048$. Note, however, that this is not the case in a *Planck*-like mission, where the noise contamination in the E mode is more important and the improvement from T only to T and polarization is 20–30 per cent (Martínez-González & Planck Collaboration 2012).

6.2 Bias

As expected, the bias (or the relative bias, shown in Figs 6 and 7) rapidly increases with ℓ_{max} . Its value is also strongly dependent on the frequency. In fact, for example in the case of no clustering, the bias is directly proportional to the amplitude of the shot-noise bispectrum of point sources (see equation 21).

If we consider only temperature and $S_c = 0.3$ Jy, the bias is practically negligible at 165 and 225 GHz (i.e. where $b_{\ell_1 \ell_2 \ell_3}^{\text{ps, white}}$ is minimum) for the local, orthogonal and flat shapes. The relative bias is less than 1 also for the equilateral shape, which is the shape most affected by point sources. As for $\sigma(f_{\text{nl}})$, the bias strongly increases at 375 GHz, being typically more than one order of magnitude higher than at 225 GHz. At 105 GHz it is still small, with a relative value less than 1 (apart from the equilateral shape, with $\Delta f_{\text{nl}} = 107$). At lower frequencies, however, $\Delta f_{\text{nl}} \gg \sigma(f_{\text{nl}})$.

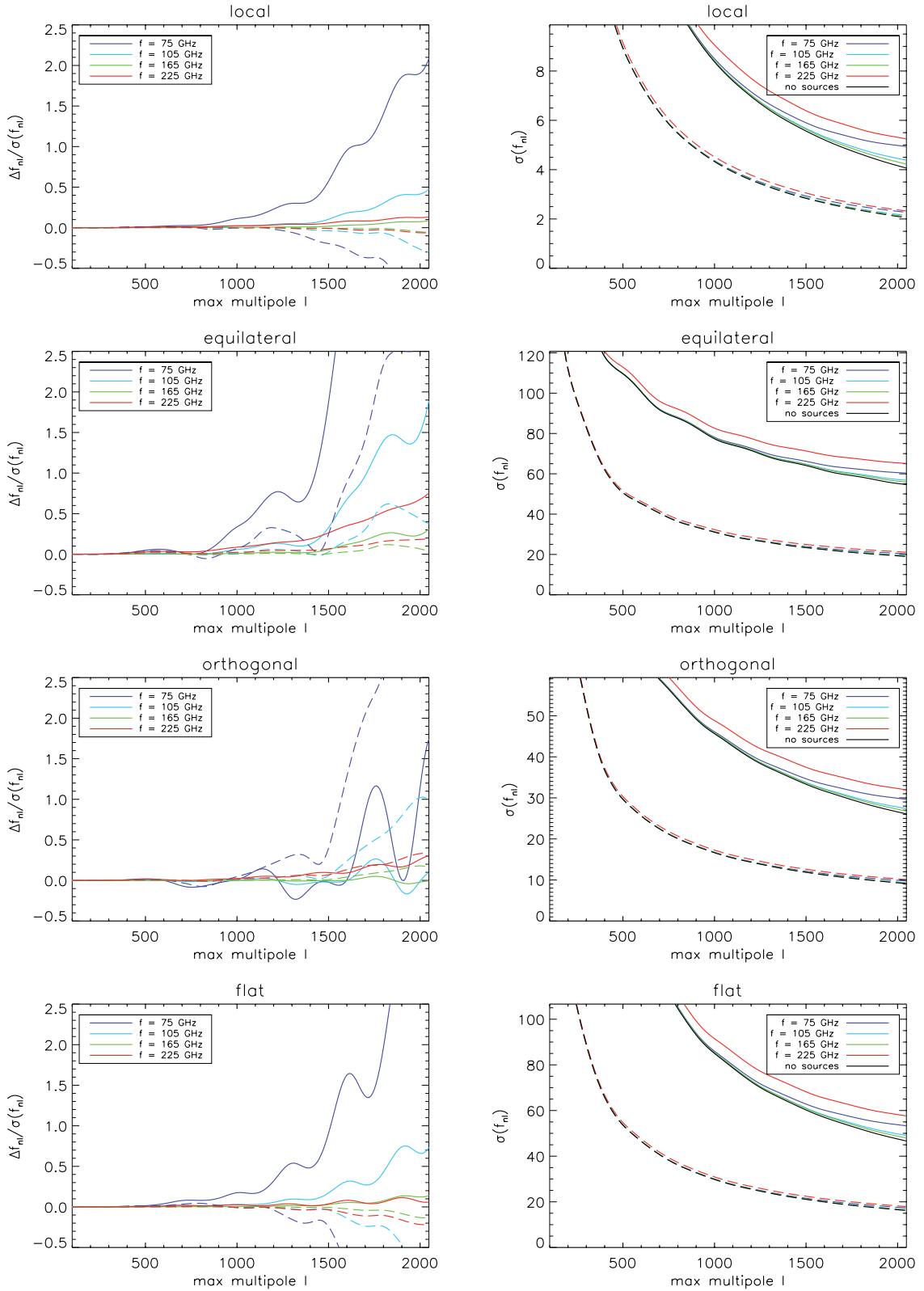


Figure 6. From left to right: the relative bias $\Delta(f_{\text{nl}})/\sigma(f_{\text{nl}})$ and the f_{nl} uncertainty $\sigma(f_{\text{nl}})$ produced by unresolved point sources as a function of ℓ_{max} for the ideal case. We consider the flux density limit $S_c = 0.3$ Jy and the following frequencies: 75, 105, 165, 225 GHz and the configuration without point sources (black line in right-hand panels). From top to bottom, we plot results for the local, equilateral, orthogonal and flat f_{nl} shapes. The solid lines take into account only temperature, while the long dashed lines include temperature and polarization.

Table 2. The expected uncertainty $\sigma(f_{\text{nl}})$, bias Δf_{nl} and relative bias $\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$ at $\ell_{\text{max}} = 2048$ from temperature only and temperature plus polarization for an ideal mission. $S_c = 0.3$ Jy is adopted.

	Case	T	T	T	T	T	T	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$
	Freq.	45	75	105	165	225	375	45	75	105	165	225	375
Local	$\sigma(f_{\text{nl}})$	7.4	4.9	4.4	4.2	5.2	29.7	2.8	2.3	2.1	2.1	2.3	3.3
	Δf_{nl}	82.1	10.3	2.1	0.4	0.7	25.6	-9.7	-2.6	-0.6	-0.1	-0.1	-0.1
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	11.1	2.1	0.5	0.1	0.1	0.9	-3.5	-1.1	-0.3	-0.1	-0.1	-0.0
Equilateral	$\sigma(f_{\text{nl}})$	75.4	60.4	56.8	55.9	65.1	232.4	24.0	20.3	19.5	19.2	21.1	35.3
	Δf_{nl}	4629.3	599.9	107.2	17.2	49.1	950.7	480.2	54.5	7.3	0.8	4.0	17.9
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	61.4	9.9	1.9	0.3	0.8	4.1	20.0	2.7	0.4	0.0	0.2	0.5
Orthogonal	$\sigma(f_{\text{nl}})$	40.1	29.7	27.4	26.8	32.0	139.0	11.6	9.8	9.4	9.2	10.1	16.5
	Δf_{nl}	955.5	51.7	2.5	-0.2	9.7	371.4	308.7	48.3	9.4	1.6	3.4	7.2
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	23.8	1.7	0.1	-0.0	0.3	2.7	26.7	4.9	1.0	0.2	0.3	0.4
Flat	$\sigma(f_{\text{nl}})$	73.7	53.3	49.0	47.9	57.6	284.7	20.6	17.4	16.6	16.4	17.9	29.1
	Δf_{nl}	598.6	150.5	36.0	6.7	3.4	-66.0	-311.6	-56.3	-12.2	-2.2	-3.9	-5.2
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	8.1	2.8	0.7	0.1	0.1	-0.2	-15.1	-3.2	-0.7	-0.1	-0.2	-0.2

Table 3. The expected uncertainty $\sigma(f_{\text{nl}})$, bias Δf_{nl} and relative bias $\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$ at $\ell_{\text{max}} = 2048$ from temperature only and temperature plus polarization for an ideal mission. $S_c = 1$ Jy is adopted.

	Case	T	T	T	T	T	T	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$
	Freq.	45	75	105	165	225	375	45	75	105	165	225	375
Local	$\sigma(f_{\text{nl}})$	9.9	5.9	4.9	4.4	5.3	29.8	3.2	2.5	2.3	2.1	2.3	3.3
	Δf_{nl}	351.1	68.5	17.6	3.4	2.3	25.6	-5.9	-12.6	-4.4	-1.0	-0.5	-0.1
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	35.6	11.6	3.6	0.8	0.4	0.9	-1.9	-5.1	-2.0	-0.5	-0.2	-0.0
Equilateral	$\sigma(f_{\text{nl}})$	88.7	66.6	59.9	57.0	65.6	229.8	27.7	21.8	20.2	19.5	21.2	35.3
	Δf_{nl}	17689.6	4075.1	1017.5	174.2	135.5	948.9	2226.3	390.7	90.7	11.5	10.8	20.5
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	199.4	61.2	17.0	3.1	2.1	4.1	80.4	17.9	4.5	0.6	0.5	0.6
Orthogonal	$\sigma(f_{\text{nl}})$	50.2	33.8	29.3	27.5	32.3	138.6	13.4	10.5	9.7	9.4	10.1	16.5
	Δf_{nl}	5411.3	573.7	80.9	3.6	16.4	377.8	1007.2	297.6	82.1	15.1	9.3	7.1
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	107.8	17.0	2.8	0.1	0.5	2.7	75.3	28.4	8.5	1.6	0.9	0.4
Flat	$\sigma(f_{\text{nl}})$	94.6	61.3	52.7	49.2	58.2	284.5	23.8	18.6	17.3	16.6	18.0	29.2
	Δf_{nl}	455.7	784.2	263.4	59.1	26.6	-69.0	-770.4	-326.8	-96.4	-19.6	-10.8	-4.1
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	4.8	12.8	5.0	1.2	0.5	-0.2	-32.4	-17.5	-5.6	-1.2	-0.6	-0.1

Moreover, we find that Δf_{nl} for the local, equilateral and flat shapes is negative, although negligible, at low multipoles ($\ell \leq 200$) and it becomes positive at higher multipoles.⁹ The bias for the orthogonal shape is oscillating around zero and therefore, in this case, there are certain multipoles where $\Delta f_{\text{nl}} \simeq 0$.

It is interesting to notice that when instrumental noise and beam function are introduced, the bias (and the relative bias) are equal or even lower than in the ideal case at $\ell_{\text{max}} = 2048$. One simple explanation is that the beam and the noise prevent the high- ℓ terms in the sum of equations (21) and (24) to diverge.

When the polarization is considered, the bias suffers important variations for each different shape in its amplitude and its sign. This is dependent on the extra terms in the estimator: the E -mode primordial bispectrum, the E -mode point source bispectrum and also the inverse of the covariance matrix in equation (26). The bias of the local and flat shapes becomes negative, whereas in the equilateral and orthogonal shapes it keeps a positive value. The local and flat shapes bias are also reduced in amplitude but the strongest improvement is for the equilateral shape. On the contrary,

at certain frequencies the amplitude of the bias of the orthogonal shape increases compared with the case of temperature only (see Tables 2 and 3). Finally, it can be noticed that the amplitude of the bias is typically lower in a *COre*-like mission than in the ideal case, as effect of considering instrumental properties in equations (21) and (24).

7 CONCLUSIONS

In this paper we present a full set of forecasts on the contamination produced by unresolved extragalactic point sources on the primordial non-Gaussianity f_{nl} parameter – hopefully detectable by the analysis of future high-resolution CMB (all-sky and low-noise) anisotropy maps – in the frequency interval 45–375 GHz. We have considered two scenarios: (a) an ideal case without instrumental noise nor instrumental beams; (b) a future experiment with *COre*-like instrumental parameters. We have characterized all source populations relevant in this frequency interval by using two recently published models: the T11 model for extragalactic radio sources and the L11 model for far-IR/submm sources. It is important to remind here that these two models have proved very successful in the explanation of all of the most recent and published results on number counts (and on other relevant statistics) of extragalactic point sources in the frequency range from ~ 30 to ~ 1000 GHz.

⁹This change in the sign of the ‘local’ bias is not found in Lacasa et al. (2012), where the bias due to IR sources is still negative at $\ell = 700$ and 2048. This different result may be dependent on the actual details of the characterization of the IR sources.

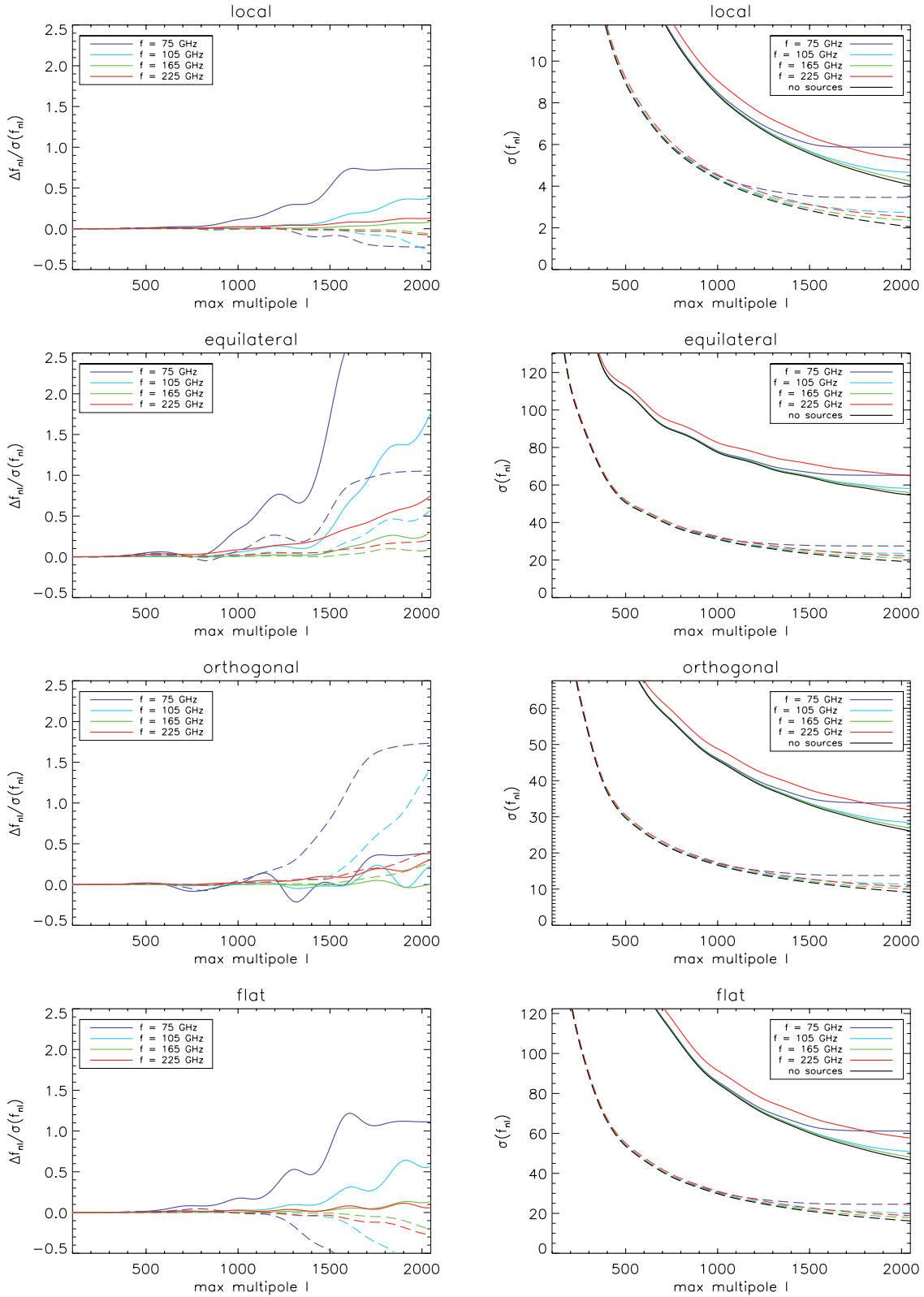


Figure 7. As in Fig. 6 but for a CORe-like mission.

As a first step, the angular power spectrum and bispectrum due to unresolved extragalactic point sources for both temperature and polarization have been estimated and compared. According to the above quoted models, the contaminating signal produced

by far-IR/submm selected extragalactic sources starts to dominate over the ERS one at $\nu \gtrsim 200$ GHz in the temperature (or total intensity) angular power spectrum and bispectrum. In spite of their very low (average) level of polarized emission, this source

Table 4. The expected uncertainty $\sigma(f_{\text{nl}})$, bias Δf_{nl} and relative bias $\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$ at $\ell_{\text{max}} = 2048$ from temperature only and temperature plus polarization for a *COrE*-like mission. $S_c = 0.3$ Jy is adopted.

	Case	T	T	T	T	T	T	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$
	Freq.	45	75	105	165	225	375	45	75	105	165	225	375
Local	$\sigma(f_{\text{nl}})$	9.7	5.9	4.7	4.2	5.3	29.8	6.0	3.5	2.7	2.3	2.5	20.6
	Δf_{nl}	16.0	4.3	1.8	0.4	0.7	25.1	1.9	-0.8	-0.7	-0.1	-0.2	8.2
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	1.7	0.7	0.4	0.1	0.1	0.8	0.3	-0.2	-0.3	-0.1	-0.1	0.4
Equilateral	$\sigma(f_{\text{nl}})$	84.3	65.2	58.1	56.0	65.2	232.6	41.1	27.5	23.5	21.0	22.2	147.2
	Δf_{nl}	421.9	202.4	102.7	17.5	49.0	935.9	66.2	28.9	13.5	1.9	4.6	343.8
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	5.0	3.1	1.8	0.3	0.8	4.0	1.6	1.1	0.6	0.1	0.2	2.3
Orthogonal	$\sigma(f_{\text{nl}})$	49.7	33.8	28.4	26.9	32.0	139.1	21.7	13.8	11.4	10.1	10.6	82.0
	Δf_{nl}	39.9	12.9	6.5	-0.1	9.7	367.8	12.1	23.8	16.1	2.8	4.4	130.2
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	0.8	0.4	0.2	-0.0	0.3	2.6	0.6	1.7	1.4	0.3	0.4	1.6
Flat	$\sigma(f_{\text{nl}})$	93.8	61.2	50.9	48.1	57.7	285.1	39.1	24.5	20.2	17.9	18.9	148.1
	Δf_{nl}	190.3	68.1	29.0	6.5	3.4	-68.8	10.4	-26.3	-20.4	-3.7	-5.2	-38.4
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	2.0	1.1	0.6	0.1	0.1	-0.2	0.3	-1.1	-1.0	-0.2	-0.3	-0.3

Table 5. The expected uncertainty $\sigma(f_{\text{nl}})$, bias Δf_{nl} and relative bias $\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$ at $\ell_{\text{max}} = 2048$ from temperature only and temperature plus polarization for *COrE*-like mission. $S_c = 1$ Jy is adopted.

	Case	T	T	T	T	T	T	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$	$T+E$
	Freq.	45	75	105	165	225	375	45	75	105	165	225	375
Local	$\sigma(f_{\text{nl}})$	11.1	6.4	5.0	4.4	5.3	29.9	6.6	3.7	2.9	2.4	2.5	20.6
	Δf_{nl}	130.7	37.1	15.6	3.4	2.3	25.0	15.3	-7.0	-6.0	-1.2	-0.6	8.1
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	11.8	5.8	3.1	0.8	0.4	0.8	2.3	-1.9	-2.1	-0.5	-0.2	0.4
Equilateral	$\sigma(f_{\text{nl}})$	92.4	68.9	60.6	57.1	65.7	230.0	43.7	28.4	24.0	21.2	22.3	146.3
	Δf_{nl}	3428.4	1731.9	916.2	175.0	135.3	933.6	515.4	235.0	117.5	19.5	12.6	346.7
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	37.1	25.1	15.1	3.1	2.1	4.1	11.8	8.3	4.9	0.9	0.6	2.4
Orthogonal	$\sigma(f_{\text{nl}})$	55.3	36.2	30.0	27.5	32.3	138.7	23.3	14.3	11.7	10.2	10.7	81.8
	Δf_{nl}	553.1	161.5	79.9	4.4	16.4	374.0	110.3	202.6	136.9	25.4	12.2	132.5
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	10.0	4.5	2.7	0.2	0.5	2.7	4.7	14.2	11.7	2.5	1.1	1.6
Flat	$\sigma(f_{\text{nl}})$	105.8	65.8	53.9	49.3	58.3	284.9	42.0	25.5	20.8	18.1	19.0	147.7
	Δf_{nl}	1232.1	523.1	232.3	58.3	26.5	-72.1	59.4	-227.2	-172.4	-33.0	-14.7	-39.3
	$\Delta f_{\text{nl}}/\sigma(f_{\text{nl}})$	11.6	7.9	4.3	1.2	0.5	-0.3	1.4	-8.9	-8.3	-1.8	-0.8	-0.3

population is also relevant in polarization, especially at large scales, due to the enhanced signal produced by the strong clustering of high- z far-IR/submm spheroidal galaxies. On the contrary, their effect is completely negligible at frequencies $\lesssim 100$ GHz. On the other hand, unresolved ERS constitute a more relevant contaminant in all the frequency range considered here and, thus, their effect on the CMB anisotropy signal has to be properly modelled, for a correct evaluation of their impact. In any case, their signal is found to be dominant over the intrinsic CMB one only at frequencies below 75 GHz, thus out of the CMB cosmological window.

We have studied the expected impact of undetected point sources on the f_{nl} parameter in terms of uncertainty and bias, for the local, equilateral, orthogonal and flat shapes. For a *COrE*-like mission, we find out that the best frequency for a f_{nl} analysis is ~ 165 GHz. At this frequency (and $S_c = 0.3$ Jy), the effect of unresolved point sources on the accuracy in the determination of f_{nl} is almost negligible. The uncertainty on f_{nl} does not significantly change whether or not their contribution is included, and estimated values of f_{nl} are essentially unbiased (apart from the equilateral shape for which $\Delta f_{\text{nl}}/\sigma(f_{\text{nl}}) \sim 0.3$ at $\ell_{\text{max}} = 2048$). At the frequencies 75 and 105 GHz, results are similar to 165-GHz ones for the $\sigma(f_{\text{nl}})$, but f_{nl} can be significantly biased, mainly for the equilateral and the flat shape; on the other hand, at 225 GHz we observe an increase in the f_{nl} error-bars. If we move far from this frequency range the

capability of an accurate f_{nl} measurement fast degrades, especially if only temperature is used.

When polarization is considered, error-bars in the f_{nl} estimate are reduced by at least a factor of 2 for the ideal case (in agreement with estimates in literature; see e.g. Babich et al. 2004) and they are not affected by point sources in all the frequency range 75–225 GHz. Moreover, polarization can also help to significantly reduce the bias induced by unresolved sources. As example, for the equilateral shape, i.e. the shape most affected by point sources, the bias is a factor from 3 to 10 lower than from temperature only analysis. We can say therefore that polarization allows to expand the best range of frequencies where f_{nl} can be safely determined.

Discrepancies in the far-IR source bispectrum are observed when compared with Lacasa et al. (2012). They are attributed to the different ways to model the far-IR sources (see Section 5). This dependency might affect the bias for the considered shapes especially at high frequencies (see Section 6).

Finally, as expected, the bias induced by point sources is strongly related to the flux density limit down to which sources are detected and removed/masked. Our results seem to point out that $S_c = 0.3$ Jy is low enough to have unbiased f_{nl} values at least at 165 and 225 GHz for a *COrE*-like mission. At different frequencies, lower flux density limits are required. On the other hand, f_{nl} uncertainty is weakly affected by the choice of S_c because it is dominated by the intrinsic f_{nl} variance (i.e. independent of contaminants), that decreases when

increasing ℓ_{\max} . We want to stress the importance of high-resolution (i.e. FWHM < 10 arcmin) and polarization-sensitive instruments, in order to have significant improvements in the precision of primordial f_{nl} in future experiments beyond *Planck*.

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